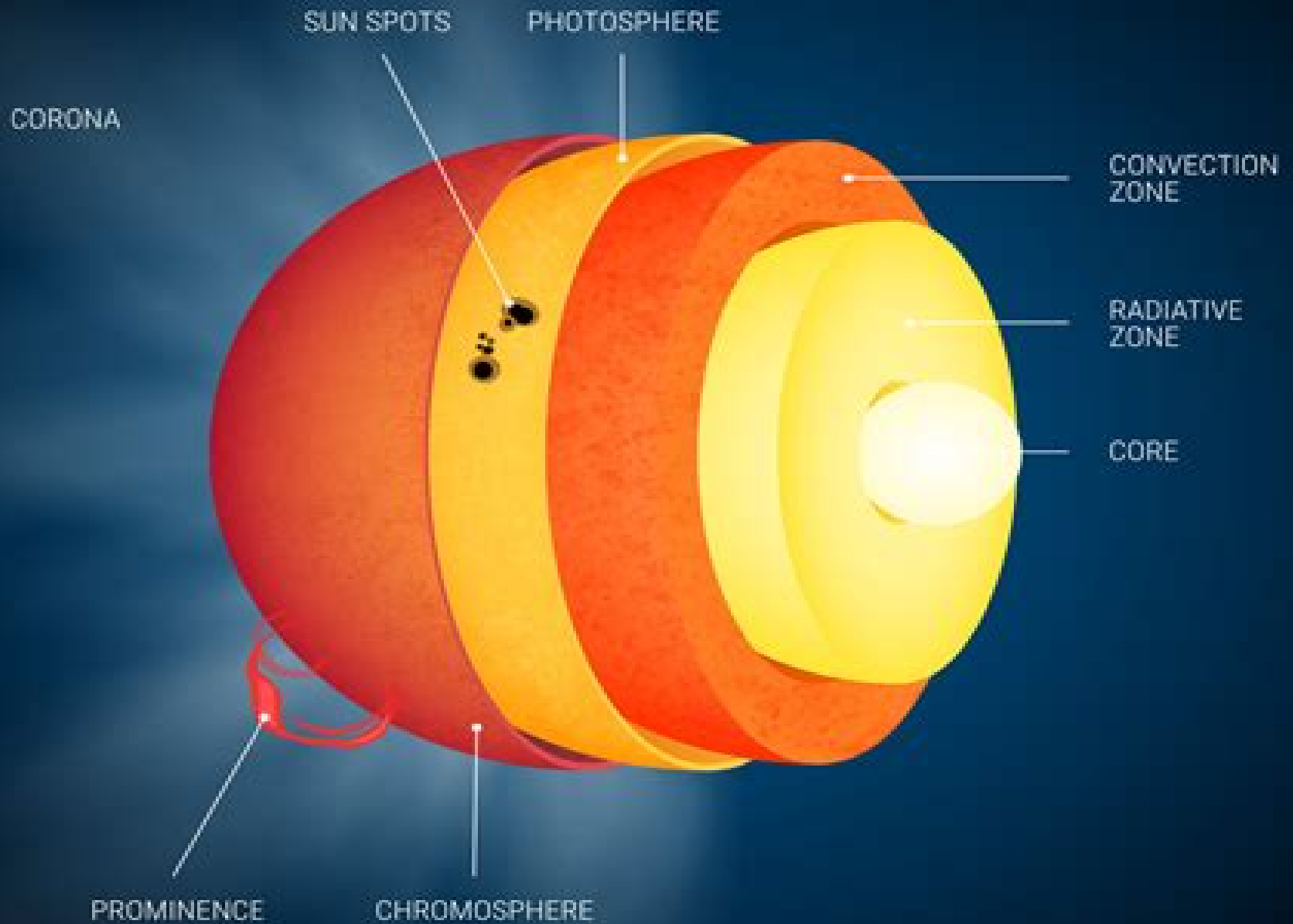




Structure Of The Quiet Photosphere The

C. De Jager



Structure Of The Quiet Photosphere The:

The Structure of the Quiet Photosphere and the Low Chromosphere C. de Jager, 2012-12-06 From 17 to 21 April 1967 a Study Week was held in the hotel De Bilderberg near Arnhem Holland with the purpose to establish a new and if possible generally acceptable working model for the quiet parts of the solar photosphere and low chromosphere The organizers of the conference hoped that even if this latter goal appeared too far to be reached such a meeting would still be useful if only for enumerating the crucial problems in solar photospheric research and for defining future subjects of research About twenty solar physicists from outside the Netherlands participated in the Study Week while some others though prevented from actively attending submitted their comments before the meeting The two above mentioned goals were reached a working model could be established yet it became clear that not everyone would agree about this model and it became obvious too that future research is strongly needed in particular in the field of line formation coherence or non coherence local thermal equilibrium while also the motion field of the photosphere and chromosphere is insufficiently known and its influence on the formation of spectral lines hardly understood

The Structure of the Quiet Photosphere and the Low Chromosphere C. De Jager, 1968-07-31 **The Structure of the Quiet Photosphere and the Low Chromosphere**, 1968 **Solar Photosphere: Structure, Convection, and Magnetic Fields** Jan Olof Stenflo, 2012-12-06

Solar and stellar photospheres constitute the layers most accessible to observations forming the interface between the interior and the outside of the stars The solar atmosphere is a rich physics laboratory in which the whole spectrum of radiative dynamical and magnetic processes that transfer energy into space can be observed As the fundamental processes take place on very small spatial scales we need high resolution observations to explore them On the other hand the small scale processes act together to form global properties of the sun which have their origins in the solar interior The rapid advances in observational techniques and theoretical modelling over the past decade made it very timely to bring together scientists from east and west to the first IAU Symposium on this topic The physics of the photosphere involves complicated interactions between magnetic fields convection waves and radiation During the past decade our understanding of these generally small scale structures and processes has been dramatically advanced New instruments on ground and in space have given us new means to study the granular convection Diagnostic methods in Stokes polarimetry have allowed us to go beyond the limitations of spatial resolution to explore the structure and dynamics of the subarcsec magnetic structures Extensive numerical simulations of the interaction between convection and magnetic fields using powerful supercomputers are providing deepened physical insight Granulation magnetic fields and dynamo processes are being explored in the photospheres of other stars guided by our improved understanding of the solar photosphere

The Structure of the quiet photosphere and the low chromosphere Cornelis de Jager, 1968 **The Origin and Dynamics of Solar Magnetism**

M.J. Thompson, A. Balogh, J.L. Culhane, Å. Nordlund, S.K. Solanki, J.-P. Zahn, 2009-05-01 Starting in 1995 numerical modeling of

the Earth's dynamo has flourished with remarkable success. Direct numerical simulation of convection driven MHD flow in a rotating spherical shell shows magnetic fields that resemble the geomagnetic field in many respects: they are dominated by the axial dipole of approximately the right strength; they show spatial power spectra similar to that of Earth; and the magnetic field morphology and the temporal variation of the field resembles that of the geomagnetic field (Christensen and Wicht 2007). Some models show stochastic dipole reversals whose details agree with what has been inferred from paleomagnetic data (Glatzmaier and Roberts 1995; Kutzner and Christensen 2002; Wicht 2005). While these models represent direct numerical simulations of the fundamental MHD equations without parameterized induction effects, they do not match actual planetary conditions in a number of respects. Specifically, they rotate too slowly, are much less turbulent, and use a viscosity and thermal diffusivity that is far too large in comparison to magnetic diffusivity. Because of these discrepancies, the success of geodynamo models may seem surprising. In order to better understand the extent to which the models are applicable to planetary dynamos, scaling laws that relate basic properties of the dynamo to the fundamental control parameters play an important role. In recent years, first attempts have been made to derive such scaling laws from a set of numerical simulations that span the accessible parameter space (Christensen and Tilgner 2004; Christensen and Aubert 2006).

The Pulsations of the Sun and the Stars Jean-Pierre Rozelot, Coralie Neiner, 2011-07-03. This volume of lecture notes brings together the knowledge on pulsations of the Sun and the stars with a particular emphasis on recent observations and modelling and on the influence of pulsations of other physical processes. The book begins with an extensive introduction to helioseismology. The solar cycle and gravity modes are discussed before the focus is widened from helioseismology to asteroseismology, which is detailed in a series of specific chapters. Based on courses given at a graduate school, these tutorial lecture notes will be of interest and useful to a rather broad audience of scientists and students.

Progress in Solar Physics C. de Jager, Zdenek Svestka, 2012-12-06. **THE STRUCTURE OF THE QUIET PHOTOSPHERE AND THE LOW CHROMOSPHERE-PROCEEDINGS OF THE BILDERBERG CONFERENCE.** , **The Stars** Evry L. Schatzman, Francoise Praderie, 1993-04-29. Following an introductory chapter on stellar fundamentals, the story begins of the lives and deaths of the stars. The authors take us on a journey from the sun, a comparatively young star, to supernovae, manifestations of dramatic death. On the way, the reader, advanced undergraduate or beginning postgraduate, is presented with such topics as stellar evolution, the hydrodynamics of stellar interiors, variability, and solar and stellar activity. This comprehensive but rigorous text, building on the theoretical and observational advances of recent years, for instance, astrophysical theory in the light of Supernova 1987A, is required reading for all serious students of astronomy.

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